

Geodesic motion in the space-time of a non-compact boson star

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We study the geodesic motion of test particles in the space-time of non-compact boson stars. These objects are made of a self-interacting scalar field and – depending on the scalar field’s mass – can be as dense as neutron stars or even black holes. In contrast to the former these objects do not contain a well-defined surface, while in contrast to the latter the space-time of boson stars is globally regular, can – however – only be given numerically. Hence, the geodesic equation also has to be studied numerically. We discuss the possible orbits for massive and massless test particles and classify them according to the particle’s energy and angular momentum. The space-time of a boson star approaches the Schwarzschild space-time asymptotically, however deviates strongly from it close to the center of the star. As a consequence, we find additional bound orbits of massive test particles close to the center of the star that are not present in the Schwarzschild case. Our results can be used to make predictions about extreme-mass-ratio inspirals (EMRIs) and we hence compare our results to recent observational data of the stars orbiting *Sagittarius A** - the radiosource at the center of our own galaxy.

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I. INTRODUCTION

Solitons play an important rôle in many areas of physics. As classical solutions of non-linear field theories, they are localized structures with finite energy, which are globally regular. In general, one can distinguish between topological and non-topological solitons. While topological solitons [1] possess a conserved quantity, the topological charge, that stems (in most cases) from the spontaneous symmetry breaking of the theory, non-topological solitons [2, 3] have a conserved Noether charge that results from a symmetry of the Lagrangian. The standard example of non-topological solitons are Q -balls [4], which are solutions of theories with self-interacting complex scalar fields. These objects are stationary with an explicitly time-dependent phase. The conserved Noether charge Q is then related to the global phase invariance of the theory and is directly proportional to the frequency. Q can e.g. be interpreted as particle number [2]. While in standard scalar field theories, it was shown that a non-renormalizable Φ^6 -potential is necessary [5], supersymmetric extensions of the Standard Model (SM) also possess Q -ball solutions [6]. In the latter case, several scalar fields interact via complicated potentials. It was shown that cubic interaction terms that result from Yukawa couplings in the superpotential and supersymmetry (SUSY) breaking terms lead to the existence of Q -balls with non-vanishing baryon or lepton number or electric charge. These supersymmetric Q -balls have been considered as possible candidates for baryonic dark matter [7] and their astrophysical implications have been discussed [8]. In [9], these objects have been constructed numerically using the exact form of a scalar potential that results from gauge-mediated SUSY breaking. However, this potential is non-differentiable at the SUSY breaking scale. In [10] a differentiable approximation of this potential was suggested and the properties of the corresponding Q -balls have been investigated in $3 + 1$ dimensions. This was extended to $d + 1$ dimensions in [11]. Q -ball solutions with a Φ^6 -potential in $3 + 1$ dimensions have been studied in detail in [5, 12, 13].

It was realized that next to non-spinning Q -balls, which are spherically symmetric, spinning solutions exist. These are axially symmetric with energy density of toroidal shape and angular momentum $J = mQ$, where Q is the Noether charge of the solution and $m \in \mathbb{Z}$ corresponds to the winding around the z -axis. Approximated solutions of the non-linear partial differential equations were constructed in [5] by means of a truncated series in the spherical harmonics to describe the angular part of the solutions. The full partial differential equation was solved numerically in [12–14]. It was also realized in [5] that in each m -sector, parity-even ($P = +1$) and parity-odd ($P = -1$) solutions exist. Parity-even and parity-odd refers to the fact that the solution is symmetric and anti-symmetric, respectively with respect to a reflection through the x - y -plane, i.e. under $\theta \rightarrow \pi - \theta$. Complex scalar field models coupled to gravity possess so-called “boson star” solutions [15–21]. In [12, 13, 22] boson stars have been considered that have flat space-time limits in the form of Q -balls. These boson stars are hence self-gravitating Q -balls. In [11, 23] the gravitating generalizations of the supersymmetric Q -balls studied in [10] have been discussed in $d + 1$ dimensions. It was found that the behaviour of the mass and charge at the critical value of the frequency depends crucially on the number of dimensions d . While in most models considered, the scalar field function is exponentially decaying and hence different notions of a boson star radius exist, this is different in models with a V-shaped potential [24, 25]. In this case, compact boson stars with a well-defined outer radius (very similar to those of “standard stars”) can be given [26–28].

In this paper we are interested in the possibility to detect boson stars through the motion of massive and massless test particles in their space-time. In particular, we will be interested in the difference between test particle motion in a boson star space-time and a Schwarzschild space-time. Since boson stars are very compact, they have been considered as alternatives to supermassive black holes [29] residing e.g. in the center of galaxies and geodesic motion of massive test particles describing extreme-mass-ratio inspirals (EMRIs) has been discussed in some particular cases in [30]. While objects with a well-defined surface as alternatives to the supermassive black hole at the center of our own galaxy, the Milky Way have been ruled out [31], the boson stars studied in this paper have a scalar field falling off exponentially at infinity and hence strictly speaking do not have a well-defined outer surface outside which the energy density and pressure, respectively, vanishes. We hence study *non-compact* boson stars in this paper and we will make a detailed analysis of the motion of massive and massless test particles in the space-time of such a boson star.

Our paper is organised as follows: in Section II, we give the field theoretical model, the Ansatz and the equations of motion to describe the space-time of a non-spinning, non-compact boson star. In Section III we discuss the geodesic equation and give our numerical results in Section IV. We conclude in Section V.

II. THE SPACE-TIME OF A NON-COMPACT BOSON STAR

In the following we will discuss the field theoretical model to describe the space-time of a non-spinning, non-compact boson star in which the test particles will move on geodesics.

The action S of the field theoretical model reads:

$$S = \int \sqrt{-g} d^4x \left(\frac{R}{16\pi G} + \mathcal{L}_m \right) \quad (1)$$

where R is the Ricci scalar, G denotes Newton's constant and the matter Lagrangian is given by

$$\mathcal{L}_m = -\partial_\mu \Phi \partial^\mu \Phi^* - U(|\Phi|) \quad (2)$$

where Φ denotes a complex scalar field and we choose as signature of the metric $(-+++)$. $U(|\Phi|)$ is the scalar field potential that arises in gauge-mediated supersymmetric breaking in the Minimal Supersymmetric extension of the Standard Model (MSSM) [9, 10]:

$$U(|\Phi|) = m^2 \eta_{\text{susy}}^2 \left(1 - \exp \left(-\frac{|\Phi|^2}{\eta_{\text{susy}}^2} \right) \right) , \quad (3)$$

where η_{susy} corresponds to the scale below which supersymmetry is broken and is roughly 1 TeV, while m denotes the scalar boson mass. For $|\Phi|^2 \ll \eta_{\text{susy}}^2$ we can give a polynomial expansion of this potential which reads

$$U(|\Phi|) = m^2 |\Phi|^2 - \frac{m^2}{\eta_{\text{susy}}^2} |\Phi|^4 + \frac{m^2}{\eta_{\text{susy}}^4} |\Phi|^6 + \mathcal{O}(|\Phi|^8) . \quad (4)$$

Neglecting terms of $|\Phi|^8$ and higher this potential has the form of the scalar potential frequently used in the study of Q -balls and boson stars [4, 5] and reads

$$\tilde{U}(|\Phi|) = \lambda |\Phi|^2 (\gamma |\Phi|^4 - \alpha |\Phi|^2 + \beta) , \quad (5)$$

where λ , γ , α and β are constants such that $\sqrt{\lambda\beta}$ corresponds to the scalar boson star mass m . In previous studies, the constants have been fixed to particular values which fulfill the requirements to find non-trivial solutions to the equations of motion.

The matter Lagrangian \mathcal{L}_m (2) is invariant under the global U(1) transformation

$$\Phi \rightarrow \Phi e^{i\chi} , \quad \chi \text{ constant} . \quad (6)$$

As such the locally conserved Noether current j^μ , $\mu = 0, 1, 2, 3$, associated to this symmetry is given by

$$j^\mu = -i (\Phi^* \partial^\mu \Phi - \Phi \partial^\mu \Phi^*) \quad \text{with} \quad j^\mu_{;\mu} = 0 . \quad (7)$$

The globally conserved Noether charge Q of the system then reads

$$Q = - \int \sqrt{-g} j^0 d^3x . \quad (8)$$

Finally, the energy-momentum tensor is given by

$$\begin{aligned} T_{\mu\nu} &= g_{\mu\nu} \mathcal{L} - 2 \frac{\partial \mathcal{L}}{\partial g^{\mu\nu}} \\ &= -g_{\mu\nu} \left[\frac{1}{2} g^{\sigma\rho} (\partial_\sigma \Phi^* \partial_\rho \Phi + \partial_\rho \Phi^* \partial_\sigma \Phi) + U(\Phi) \right] + \partial_\mu \Phi^* \partial_\nu \Phi + \partial_\nu \Phi^* \partial_\mu \Phi . \end{aligned} \quad (9)$$

The coupled system of ordinary differential equations is given by the Einstein equations

$$G_{\mu\nu} = 8\pi G T_{\mu\nu} \quad (10)$$

with $T_{\mu\nu}$ given by (9) and the Klein-Gordon equation

$$\left(\square - \frac{\partial U}{\partial |\Phi|^2} \right) \Phi = 0 . \quad (11)$$

In the following, we want to study non-spinning boson stars. For the metric we use the following Ansatz in isotropic coordinates

$$ds^2 = -f(r) dt^2 + \frac{l(r)}{f(r)} [dr^2 + r^2 d\theta^2 + r^2 \sin^2 \theta d\varphi^2] . \quad (12)$$

For the complex scalar field, we use a stationary Ansatz that contains a periodic dependence of the time-coordinate t :

$$\Phi(t, r) = e^{i\omega t} \phi(r) , \quad (13)$$

where ω is a constant and denotes the frequency.

In order to be able to use dimensionless quantities we introduce the following rescalings

$$r \rightarrow \frac{r}{m} \quad , \quad \omega \rightarrow m\omega \quad , \quad \phi \rightarrow \eta_{\text{susy}}\phi \quad (14)$$

and find that the equations depend only on the dimensionless coupling constant

$$\kappa = 8\pi G\eta_{\text{susy}}^2 = 8\pi \frac{\eta_{\text{susy}}^2}{M_{\text{pl}}^2} \quad , \quad (15)$$

where M_{pl} is the Planck mass. Note that κ denotes the ratio between the SUSY breaking scale and the Planck mass which we expect to be small $\kappa \sim 10^{-31}$ if the SUSY breaking scale is on the order of 1 TeV. In this paper, we will however also study larger values of κ in order to understand the qualitative behaviour of the solutions in curved space-time. Note that with these rescalings the scalar boson mass m becomes equal to unity.

Using the rescalings the Einstein equations read

$$f'' = -\frac{1}{2} \frac{1}{fl} \left[4\kappa fl^2 (1 - \exp(-\phi^2)) - 8\kappa\omega^2 l^2 \phi^2 + \frac{4}{r} flf' - 2lf'^2 + fl'f' \right] \quad (16)$$

and

$$l'' = -\frac{1}{2} \frac{1}{fl} \left[8\kappa l^3 (1 - \exp(-\phi^2)) - 8\kappa\omega^2 \frac{l^3}{f} \phi^2 + \frac{6}{r} fl l' - fl'^2 \right] \quad , \quad (17)$$

while the Klein-Gordon equation for the scalar field function is

$$\phi'' = -\frac{1}{2} \frac{1}{fl} \left[2\omega^2 \frac{l^2}{f} \phi - 2l^2 \phi \exp(-\phi^2) + \frac{4}{r} fl\phi' + fl'\phi' \right] \quad . \quad (18)$$

The prime here and in the following denotes the derivative with respect to r .

The equations (16)-(18) have to be solved subject to appropriate boundary conditions. These are given by the requirement of regularity at the origin

$$f'|_{r=0} = 0 \quad , \quad l'|_{r=0} = 0 \quad , \quad \phi'|_{r=0} = 0 \quad (19)$$

and by the requirement of finite energy, asymptotically flat solutions

$$f(r = \infty) = 1 \quad , \quad l(r = \infty) = 1 \quad , \quad \phi(r = \infty) = 0 \quad . \quad (20)$$

The set of coupled, nonlinear ordinary differential equations (16)-(18) has been studied in detail in [23], however in Schwarzschild-like coordinates. We will quote the relevant results in the following when necessary, but refer the reader to that paper for details on the field theoretical solutions.

The behaviour of the metric function f determines the mass M of the solution. This is given by [12, 13]

$$M = \frac{1}{2} \lim_{r \rightarrow \infty} r^2 \partial_r f \quad , \quad (21)$$

while the explicit expression for the Noether charge reads

$$Q = 8\pi\omega \int_0^\infty \frac{\sqrt{l}}{f} r^2 \phi^2 dr \quad . \quad (22)$$

Note that the mass M is measured in units of M_{pl}^2/m .

At spatial infinity $r \rightarrow \infty$ the scalar field function decays exponentially

$$\phi(r \gg 1) = \frac{1}{r} \exp(-\sqrt{1 - \omega^2} r) + \dots \quad (23)$$

Hence $\phi(r) \neq 0$ for $r < \infty$. As such the boson star does not possess a well-defined surface and different notions of the radius of the boson star are possible (see discussion below). While boson stars with well-defined surface do exist [28] (so-called ‘‘compact boson stars’’), which hence rather resemble astrophysical objects, boson stars without proper surface could play an important rôle as alternatives to supermassive black holes in the center of galaxies.

III. THE GEODESIC EQUATION

We consider the geodesic equation

$$\frac{d^2 x^\mu}{d\tau^2} + \Gamma_{\rho\sigma}^\mu \frac{dx^\rho}{d\tau} \frac{dx^\sigma}{d\tau} = 0 , \quad (24)$$

where $\Gamma_{\rho\sigma}^\mu$ denotes the Christoffel symbol given by

$$\Gamma_{\rho\sigma}^\mu = \frac{1}{2} g^{\mu\nu} (\partial_\rho g_{\sigma\nu} + \partial_\sigma g_{\rho\nu} - \partial_\nu g_{\rho\sigma}) \quad (25)$$

and τ is an affine parameter such that for time-like geodesics $d\tau^2 = g_{\mu\nu} dx^\mu dx^\nu$ corresponds to proper time.

The geodesic Lagrangian \mathcal{L}_g for a point particle in the space-time (12) reads

$$\mathcal{L}_g = \frac{1}{2} g_{\mu\nu} \frac{dx^\mu}{ds} \frac{dx^\nu}{ds} = -\frac{1}{2} \delta = \frac{1}{2} \left[-f \left(\frac{dt}{d\tau} \right)^2 + \frac{l}{f} \left(\frac{dr}{d\tau} \right)^2 + \frac{l}{f} r^2 \left(\frac{d\theta}{d\tau} \right)^2 + \frac{l}{f} r^2 \sin^2 \theta \left(\frac{d\varphi}{d\tau} \right)^2 \right] , \quad (26)$$

where $\delta = 0$ for massless particles and $\delta = 1$ for massive particles, respectively.

The constants of motion are the energy E and the angular momentum L of the particle (absolute value and direction). To fix the direction, we choose $\theta = \pi/2$. Hence the motion will occur in the equatorial plane. The conserved quantities then are

$$E := f \frac{dt}{d\tau} , \quad L := \frac{l}{f} r^2 \frac{d\varphi}{d\tau} . \quad (27)$$

Using these constants of motion we can rewrite (26) as follows

$$\left(\frac{dr}{d\tau} \right)^2 = \frac{1}{l} (E^2 - V_{\text{eff}}(r)) , \quad (28)$$

where $V_{\text{eff}}(r)$ is the effective potential

$$V_{\text{eff}}(r) = f \left(\delta + \frac{f L^2}{l r^2} \right) . \quad (29)$$

We then find

$$\varphi - \varphi_0 = \int \frac{dr}{\sqrt{E^2 - V_{\text{eff}}(r)}} \frac{L f}{r^2 \sqrt{l}} . \quad (30)$$

The solution to this will give $r(\varphi)$ for a test particle. Since f and l are only given numerically, $r(\varphi)$ can only be determined numerically.

IV. NUMERICAL RESULTS

We have solved the equations of motion (16)-(18) numerically using a Newton-Raphson method with an adaptive grid scheme [32]. The relative errors of the numerical integration are on the order of $10^{-10} - 10^{-13}$. The geodesic equation is integrated numerically in Fortran using the Fortran Subroutines for Mathematical Applications-IMSL MATH/LIBRARY. The accuracy of the integration as estimated from application of the method for the integration of geodesics in the Schwarzschild space-time is on the order of 10^{-5} .

A. The metric functions

In order to be able to compare the space-time of a boson star and that of a Schwarzschild black hole, we have first investigated the difference between the metric functions f and l .

κ	ω	Q	M	$\phi(0)$	$r_{\text{BS}}^{(99)}$
0.1	0.9260	101.8523	100.6465	0.4945	60.4148
0.1	0.8820	100.8126	99.7029	0.8463	59.9558
0.1	0.6800 (1)	119.6961	114.3021	2.1891	66.3377
0.1	0.6525	108.3567	106.8518	2.7331	63.0380
0.1	0.6800 (2)	86.4696	92.3413	3.2230	56.3793
1.0	0.8600	15.5600	15.1388	0.2694	80.9035
1.0	0.7620	11.1205	11.6461	0.7175	67.1487
1.0	0.8500	6.8865	8.3518	1.2718	52.0577

TABLE I: The values of ω as indicated by the black dots in Fig.2 together with the corresponding values of the charge Q , mass M and the value $\phi(0)$. The (1) and (2) denote the solutions on the 1. and 2. branch, respectively. See also [33] for further details using the potential (5).

For that, we have to remind ourselves of how the Schwarzschild metric looks like in isotropic coordinates which we use in this paper. This is given by (12) with

$$f(r) = \left(\frac{1 - \frac{M}{2r}}{1 + \frac{M}{2r}} \right)^2, \quad l(r) = \left(1 - \frac{M^2}{4r^2} \right)^2, \quad (31)$$

where M is the ADM mass of the solution and the isotropic coordinate r is related to the Schwarzschild like coordinate \tilde{r} by $\tilde{r} = r(1 + M/(2r))^2$. In Fig.1 we show the metric functions $f(r)$ and $l(r)$ for a boson star with given mass M and in comparison to that the metric functions for a Schwarzschild solution with the same M .

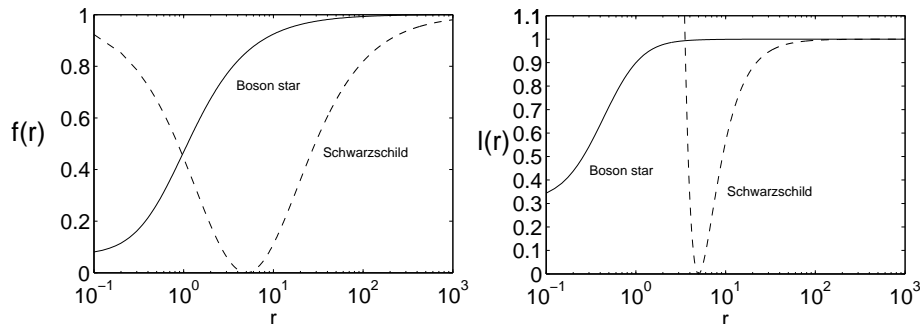


FIG. 1: We show the metric functions $f(r)$ (left) and $l(r)$ (right) for a boson star with given mass $M \approx 10$ for $\kappa = 1.0$. Also given are the metric functions $f(r)$ and $l(r)$ for a Schwarzschild solution with the same value of M . Note that the Schwarzschild solution has $f(r_h) = l(r_h) = 0$ at the event horizon $r_h = M/2$.

Clearly, the metric functions agree at large r , while for small r they deviate strongly. In particular the metric functions $f(r)$ and $l(r)$ of the Schwarzschild solution tend to zero at the event horizon $r = r_h = M/2$, while for boson stars $f(r)$ and $l(r)$ stay perfectly regular for all $r \in [0 : \infty[$. This fact will play an important rôle in the behaviour of test particles at small r , while our results already indicate here that the behaviour at large r will be similar when comparing boson star space-times with the Schwarzschild space-time.

B. Mass, charge and radius

It has been observed before that boson stars exist only in a limited range of the frequency ω . At the maximal value of ω both the mass and the charge of the boson stars tend to zero, while at the minimal value of ω a second branch of solutions starts to form and finally ends at the critical value of ω after a number of spirals. In the following, we want to discuss the motion of test particles in boson star space-times. Here we concentrate on the boson stars at some particular values of the frequency ω . These are indicated in Fig.2, where we plot the mass M as function of ω for $\kappa = 0.1$ and $\kappa = 1.0$, respectively. The exact numerical values of the mass M and charge Q are given in Table I.

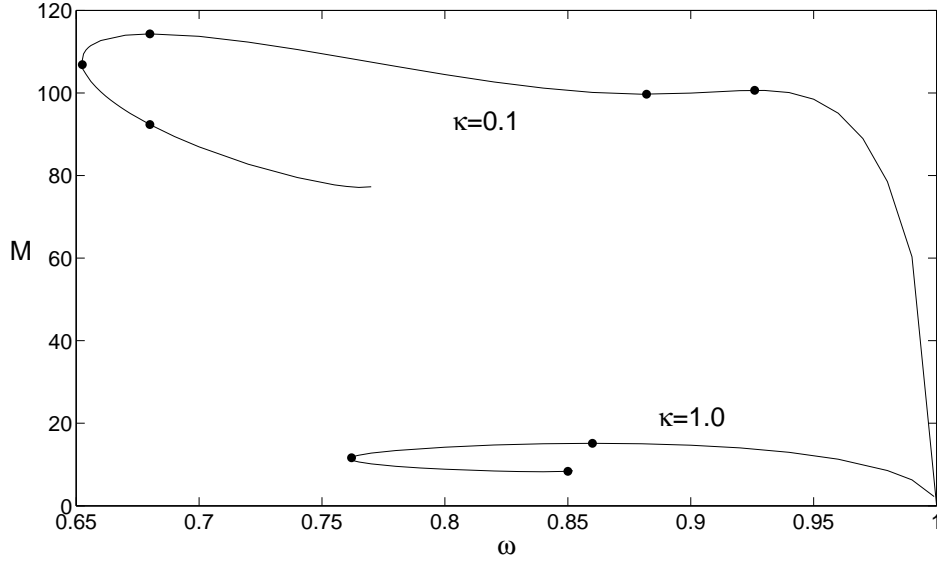


FIG. 2: The value of the mass M is shown as function of ω for two different values of κ . The black dots on the curves indicate the values of ω that we have investigated in more detail and that are given in Table I.

In the following, it will also be important to understand what the radius of the boson stars and in particular what the extend of the orbits around the boson star are. We would like to discuss this briefly in the following. Since the scalar field has an exponential decay in our model, it is difficult to define a definite radius of the boson stars discussed here. This is e.g. different in models with a V -shaped potential [28], where a definite radius of the object exists. Hence, different notions of the boson star radius can be evoked. Either, we can define the radius of the boson star as the inverse of the scalar boson star mass $r_{\text{BS}}^{(\text{m})} = m^{-1}$ (which in our rescaled variables would be equal to unity) or we could define the radius of the boson star as the radius $r_{\text{BS}}^{(99)}$ inside which 99% of the mass of the star is contained. It remains to be shown that stars consisting out of scalar fields can form. Recent results suggest that a fundamental scalar boson with mass on the order of $m \approx 125$ GeV could exist in nature [34]. If boson stars would consist of this scalar boson, their typical radii would be (in dimensionful quantities [35]) $\hat{r}_{\text{BS}}^{(\text{m})} \approx 10^{-18}$ m. If the star would consist of pions with masses of roughly 140 MeV [19], the typical radius would be $\hat{r}_{\text{BS}}^{(\text{m})} \approx 10^{-15}$ m. One could also think of much lighter scalar particles, e.g. the axion with roughly 10^{-5} eV and the dilaton with 10^{-10} eV which would correspond to much larger radii $\hat{r}_{\text{BS}}^{(\text{m})} \approx 10^{-2}$ m and $\hat{r}_{\text{BS}}^{(\text{m})} \approx 10^3$ m, respectively. When computing the radius $r_{\text{BS}}^{(99)}$ we find that this is roughly two orders of magnitude larger (see Table I for some values). On the other hand, the masses of these boson stars can be quite high. The mass M is measured in units of M_{pl}^2/m . We find that the maximal value of the mass \hat{M}_{max} is well approximated by

$$\hat{M}_{\text{max}} \approx \frac{7}{4\pi} \frac{M_{\text{pl}}^4}{\eta_{\text{SUSY}}^2 m}, \quad (32)$$

which should be compared to the typical mass of a neutron star given by $M_{\text{pl}}^3/m_{\text{neutron}}^2$ with $m_{\text{neutron}} \approx 1$ GeV.

For $m \approx 10^2$ GeV this would correspond to masses of 10^{38} GeV $\sim 10^{11}$ kg - and this inside a radius of $\hat{r}_{\text{BS}}^{(\text{m})} \approx 10^{-18}$ m (or two orders of magnitude higher if we use $r_{\text{BS}}^{(99)}$). This is an extremely dense object, much denser than a typical neutron star. However, these boson stars are stable towards gravitational collapse and due to their extremely high density have been considered as alternatives to supermassive black holes [29]. The results for radii and masses in dependence on the scalar particle mass are summarized in Table II for a dimensionless value of the boson star mass of $M = \mathcal{O}(10^2)$.

The relation between the mass M and the charge Q as given in Table I also allows us to make conclusions about the stability of the boson stars. Since Q corresponds to the particle number, i.e. the number of massive scalar bosons and these have mass $m = 1$ (in our rescaled variables), an assembly of Q massive scalar particles would have mass $M = mQ \equiv Q$. Hence, as long as the mass and charge of the boson star fulfill $M < Q$ they are stable with respect to the decay into Q individual scalar bosons. As such the boson stars on the first branch of solutions are always stable,

Particle	m	\hat{M}	$\hat{r}_{\text{BS}}^{(m)}$
Higgs	125 GeV	10^{11} kg	10^{-18} m
Pion	140 MeV	10^{14} kg	10^{-15} m
Axion	10^{-5} eV	10^{27} kg	10^{-2} m
Dilaton	10^{-10} eV	10^{32} kg	10^3 m

TABLE II: The approximate values of the mass \hat{M} and $\hat{r}_{\text{BS}}^{(m)}$ of the boson star in dimensionful units for different scalar particle with mass m and dimensionless boson star mass $M = \mathcal{O}(10^2)$. See also [33] for further details using the potential (5).

while they become unstable on parts of the second branch.

C. The effective potential and possible orbits

In all cases, we need to require that $E^2 \geq V_{\text{eff}}$ in order to have solutions to (30). At the intersection points $E^2 = V_{\text{eff}}$ we have $\frac{dr}{d\tau} = 0$. These are the turning points of the motion. Hence, the effective potential (29) can already give information about the type of orbits possible. In the following, we want to discuss which type of orbits are possible in the space-time of a non-compact boson star. We distinguish different type of orbits as follows: (a) a bound orbit is an orbit that possesses a minimal and a maximal finite radius, (b) an escape orbit is an orbit that possesses a minimal finite radius, but extends to infinity, (c) a terminating escape orbit is an orbit whose minimal radius is equal to zero and on which the particle comes from infinity to then end at the center of the boson star and (d) a terminating orbit is an orbit with finite maximal radius and minimal radius equal to zero. The possible orbits for massive test particles are summarized in Table III.

Note that in the following we will concentrate on orbits appearing at reasonably small r . Since the space-time of the boson star tends asymptotically to the Schwarzschild space-time, we have in addition all possibilities of orbits available there at large r . We will not discuss these in detail, but refer the reader to the literature (see e.g. [36, 37]).

In Fig.3 we show the effective potential V_{eff} for massive and massless test particles, respectively, for a typical boson star space-time with $\kappa = 0.1$ and $\omega \approx 0.68$ for the 1. branch as well as for the 2. branch of solutions.

Again we would like to compare this to the situation in a Schwarzschild space-time. First note that $f/l = (1 + M/2r)^{-4}$ in the Schwarzschild case such that for $r \rightarrow r_h = M/2$ the effective potential vanishes since f vanishes. As soon as the particle has crossed this horizon we cannot observe it anymore and it will eventually end up in the physical singularity at $r = 0$. For the boson star space-time on the other hand f and l are well behaved and in particular never become zero in $r \in [0 : \infty[$. As such the angular momentum term L^2/r^2 dominates the effective potential at $r \rightarrow 0$. A particle with $L \neq 0$ can never reach $r = 0$ no matter how large its energy is. Now depending on the massive test particle's energy E and angular momentum $L \neq 0$ it will either move on a bound orbit or on an escape orbit. Clearly, we can have two intersection points between E^2 and V_{eff} for massive particles as soon as the effective potential possesses a local minimum with value smaller than the asymptotic value $V_{\text{eff}}(r = \infty) = 1$ (see Fig.3(a) and Fig.3(c)). These two turning points at finite values of r correspond to the minimal radius r_{min} and the maximal radius r_{max} of the bound orbit. For both values of κ that we have studied we observe that the further we move on the branches starting at $\omega = 1$, the lower is the value of the minimum of the effective potential. This can be clearly seen when comparing the effective potential for a massive test particle with fixed L^2 for $\omega = 0.68$ on the 1. branch with that on the 2. branch. This means that we can have bound orbits for smaller values of E for a given L when moving along the branches. For massive test particles with $L \neq 0$, but E^2 large only one intersection point is possible. These particles will then move on escape orbits.

On the other hand a radially moving massive particle with $L = 0$ can reach $r = 0$ if $E^2 \geq V_{\text{eff}}(r = 0)$ which gives the condition $E^2 - f(0) > 0$. Hence, massive test particles without angular momentum and $E^2 < V_{\text{eff}}(r = \infty) = 1$ move on terminating orbits, while those with $E^2 > 1$ move on terminating escape orbits. See also Table III for the possible orbits of massive test particles.

In Fig.3(b) and Fig.3(d) we show the effective potentials for a massless test particle for different values of L in the space-time of a boson star with $\omega = 0.68$. We observe that the potential is monotonically decreasing and that no extrema of the potential exist for $L^2 > 0$, while $V_{\text{eff}}(r) \equiv 0$ for massless test particles without angular momentum, i.e. $L^2 = 0$. Hence, we can only have escape orbits for $L^2 > 0$ and terminating escape orbits for $L^2 = 0$. This is summarized in Table IV. The difference between the 1. branch solutions and 2. branch solutions does not seem to be very big for massless test particles.

For $r \gg 1$ the metric functions of both the boson star space-time and of the Schwarzschild case approach each other. Hence, the effective potential which depends only on the metric functions looks the same in both cases and in

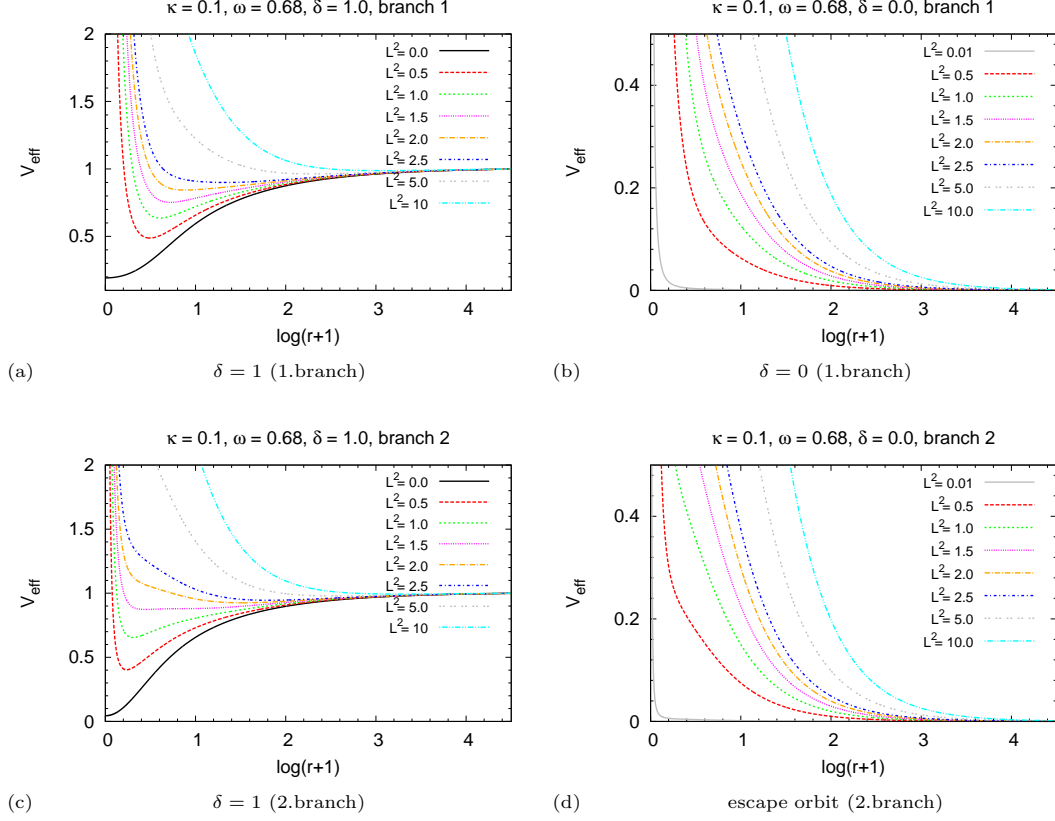


FIG. 3: We show the effective potential (29) for massive test particles (left) and massless test particles (right), respectively, for different values of L^2 . Here $\kappa = 0.1$ and $\omega \approx 0.68$. The plots on the left are for the boson star solutions of the 1. branch, the ones on the right for the boson star solutions of the 2. branch.

No. of turning points	L^2	range of r	orbit
0	$L^2 = 0$	\longrightarrow	terminating escape orbit
1	$L^2 = 0$	$\longrightarrow \bullet$	terminating orbit
2	$L^2 > 0$	$\bullet \longrightarrow \bullet$	bound orbit
1	$L^2 > 0$	$\bullet \longrightarrow$	escape orbit

TABLE III: Types of orbits possible for massive test particles moving in the space-time of a boson star.

particular $V_{\text{eff}}(r \rightarrow \infty) \rightarrow 1$ for both the boson star and the Schwarzschild space-time.

D. Examples of orbits

In the following, we want to show examples of the possible orbits mentioned above. Bound orbits of massive test particles are of particular interest for two reasons: (a) the perihelion shift of the orbits can be compared to observational data and (b) black holes are believed to form accretion discs around their centers and since boson stars have been considered as an alternative to supermassive black holes something similar should also be possible in that case. On the other hand, escape orbits of massless test particles, i.e. in particular photons can be used to compare with observational data of the light deflection.

No. of turning points	L^2	range of r	orbit
0	$L^2 = 0$	$\xrightarrow{\hspace{2cm}}$	terminating escape orbit
1	$L^2 > 0$	$\xrightarrow{\bullet \hspace{1cm}}$	escape orbit

TABLE IV: Types of orbits possible for massless test particles moving in the space-time of a boson star.

1. Massive test particles

In Fig.4 we show a bound orbit and an escape orbit of a massive test particle with $L^2 = 0.5$ in the space-time of a non-compact boson star with $\kappa = 0.1$ and the values of $\omega = 0.68$. We compare the motion of massive test particles in the space-time of the boson star solutions on the 1. and 2. branch, respectively.

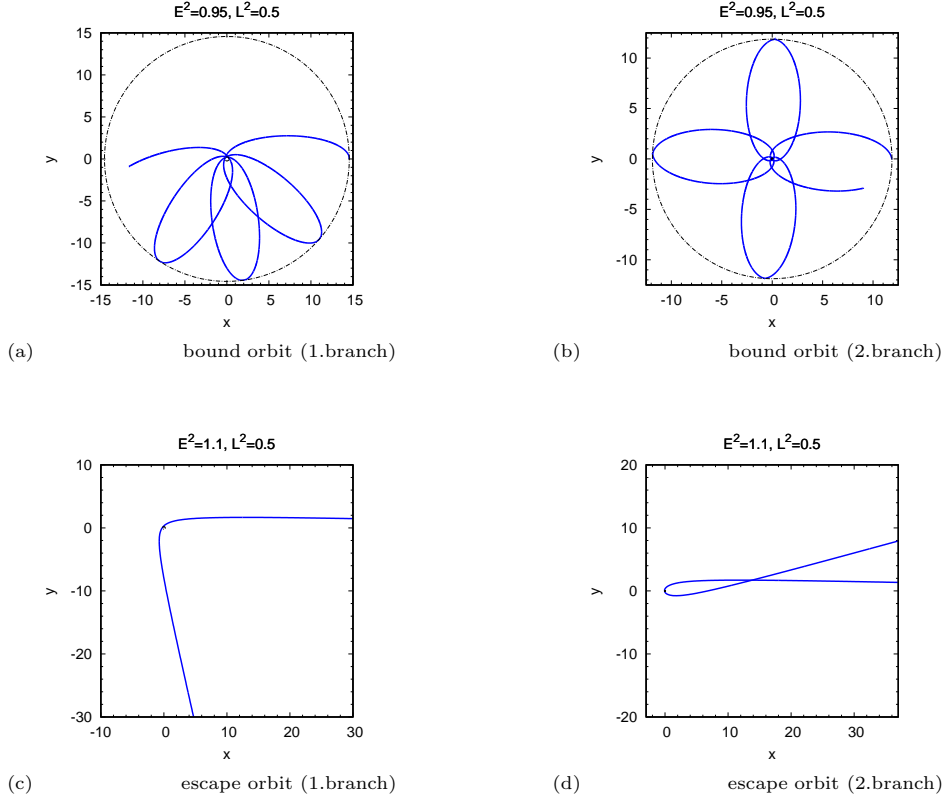


FIG. 4: We show bound as well as escape orbits for massive test particles with angular momentum $L^2 = 0.5$ in the space-time of a boson star with $\kappa = 0.1$ and $\omega = 0.68$. For bound orbits the particles have $E^2 = 0.95$, while for escape orbits we have chosen $E^2 = 1.1$. The figures on the left and right, respectively show the orbits in the space-time of the boson star solution of the 1.branch and 2.branch. In all cases the center of the boson star is located at the origin of the coordinate system.

For the bound orbits we have chosen $E^2 = 0.95 < V_{\text{eff}}(r = \infty)$, while for the escape orbits we have chosen $E^2 = 1.1 > V_{\text{eff}}(r = \infty)$. Though the frequency is the same in both cases, we see that there are differences in the particle motion. The bound orbits differ clearly in their perihelion shift. The perihelion shift in the space-time of the boson star of the 2. branch is much larger than for that of the boson star on the 1. branch (see also discussion below). The maximal radius of the orbit in the case of the 1. branch on the other hand is slightly bigger (see Table V for exact numerical values). Moreover, we observe that the further we move on the branches, the lower is the value of r_{min} , while for $\kappa = 1.0$ the value of r_{max} decreases when moving along the branches. For $\kappa = 0.1$ there seems to be a connection between the value of r_{max} and the value of M with r_{max} being largest for the maximal possible value of M (see Fig.2).

The escape orbits also differ in the sense that in the case of the 1. branch the test particles approach the boson star

E^2	L^2	κ	ω	r_{\min}	r_{\max}	e
0.97	1.0	0.1	0.9260	1.7688	19.5923	0.9959
0.95	1.5	0.1	0.8820	1.8418	11.3479	0.9867
0.95	0.5	0.1	0.6800 (1)	0.2376	14.5724	0.9998
0.95	1.5	0.1	0.6525	0.3106	12.3623	0.9997
0.95	0.5	0.1	0.6800 (2)	0.0823	11.8867	0.9999
0.97	1.0	1.0	0.8600	0.8302	27.6181	0.9995
0.97	1.0	1.0	0.7620	0.2520	22.3061	0.9999
0.95	1.0	1.0	0.8500	0.5855	9.8599	0.9982

TABLE V: The values of r_{\min} , r_{\max} as well as of the eccentricity $e := \sqrt{(r_{\max}^2 - r_{\min}^2)/r_{\max}^2}$ of bound orbits with given E^2 and L^2 in boson star space-times with κ and ω . The L^2 value is chosen close to the upper boundary up to which bound orbits are possible such that the eccentricity of the orbits is close to its maximal possible value.

and get deflected into a direction nearly perpendicular to the original one. In the case of the 2. branch the particle encircles the boson star and gets deflected into a direction similar to the one that it came from.

a. Innermost stable circular orbits (ISCOs) If the effective potential possesses a local minimum and we choose E^2 such that it is equal to the value of the effective potential at this minimum the orbits are circular orbits, so-called “innermost stable circular orbits” (ISCOs). Examples of these orbits are shown in Fig.5 for different values of ω and L^2 . Note that as soon as we have fixed the L^2 the ISCO exists only for one specific value of E^2 .

We observe that the further we move on the branches, the smaller is the radius of the ISCO of a particle of given angular momentum. This can be seen for a test particle with $L^2 = 1.5$ when comparing Fig.5(b) and Fig.5(e) as well as for a particle with $L^2 = 0.5$ when comparing Fig.5(c) and Fig.5(d). Apparently, the extend of the orbit is not only connected to the mass of the boson star since the ISCO of a $L^2 = 1.5$ particle has larger radius in the space-time of a boson star with $\omega = 0.8820$, $M = 99.7029$ as compared to that in a boson star space-time with $\omega = 0.6525$, $M = 106.8518$, while the ISCO for a $L^2 = 0.5$ particle has a larger radius in a boson star space-time with $\omega = 0.6800$, $M = 114.3021$ as compared to that in a boson star space-time with $\omega = 0.6800$, $M = 92.3413$. The message that one should take away here is that the non-trivial scalar field seems to play an important rôle in the particle motion and that no direct relation between the mass of the object and the ISCO exists as one would expect from a Newtonian treatment of the orbits.

2. Massless test particles

In Fig.6 we show the escape orbits of massless test particles with $E^2 = 0.1$, $L^2 = 1.0$ in the space-time of the boson stars given in Table I. Depending on the choice of ω and κ we find that the particles get deflected into a direction between roughly $\pi/4$ and $\pi/2$ and that for κ large enough the minimal radius of the escape orbit lies inside the Schwarzschild radius of the corresponding black hole with the same mass. We do not plot terminating escape orbits of massless test particles here since they would just correspond to a line from infinity to $r = 0$.

E. Observations of boson stars

Next to the possibility that boson stars could act as a toy model for neutron stars, boson stars have been suggested as alternatives to supermassive black holes residing in the center of galaxies [29]. As mentioned above, objects with a well-defined surface as alternatives to the supermassive black hole at the center of our own Milky Way have been ruled out [31]. We have, however argued above that our boson stars are non-compact in the sense that there is no well-defined surface beyond which the energy density and pressure, respectively, of the boson star strictly vanishes.

1. Boson star at the center of our galaxy

The center of our Milky Way possesses a very compact astronomical radio source named *Sagittarius A**. It is believed to be a supermassive black hole as observations of the motion of stars around the galactic center suggest that the mass of the central object is on the order of $\bar{M} \approx 10^{37} kg$ which corresponds to roughly $4 \cdot 10^6$ solar masses

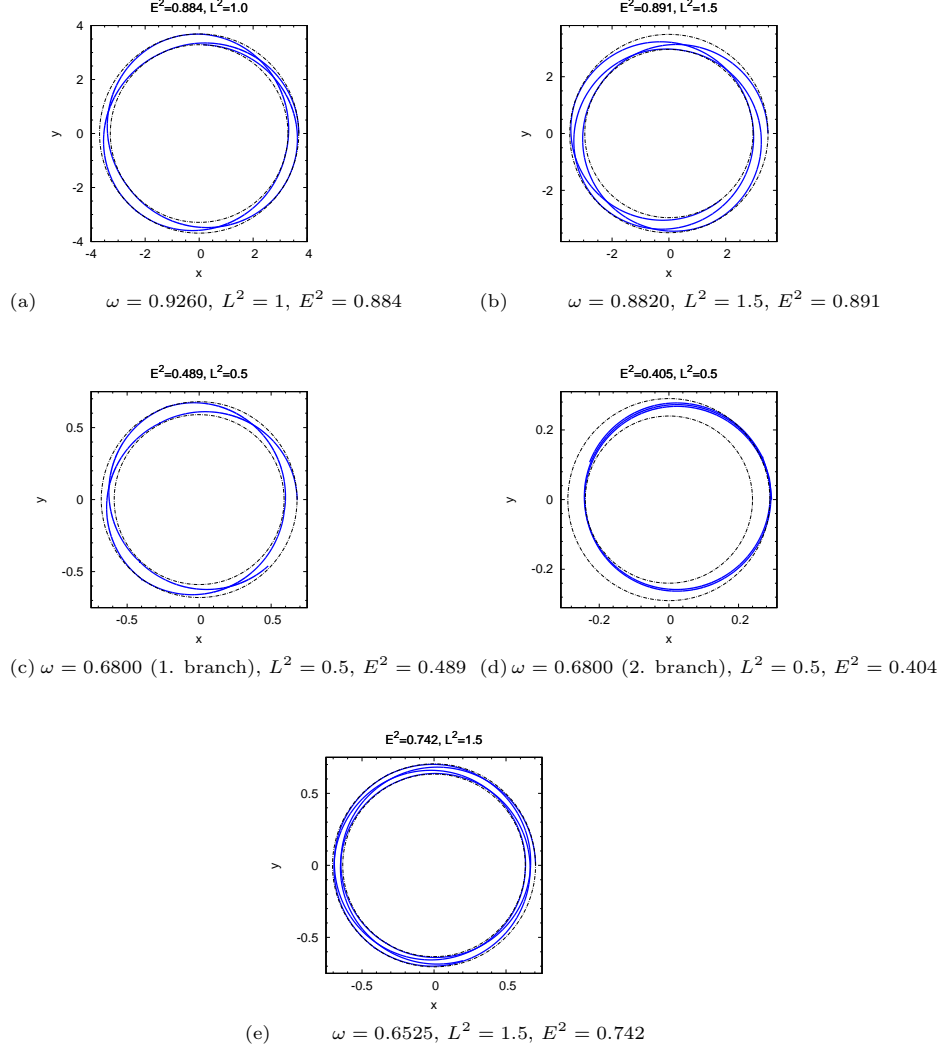


FIG. 5: We show nearly circular, so-called “innermost stable circular orbits” (ISCOs) for massive test particles with given angular momentum L^2 in the space-time of a boson star with $\kappa = 0.1$ and different values of ω .

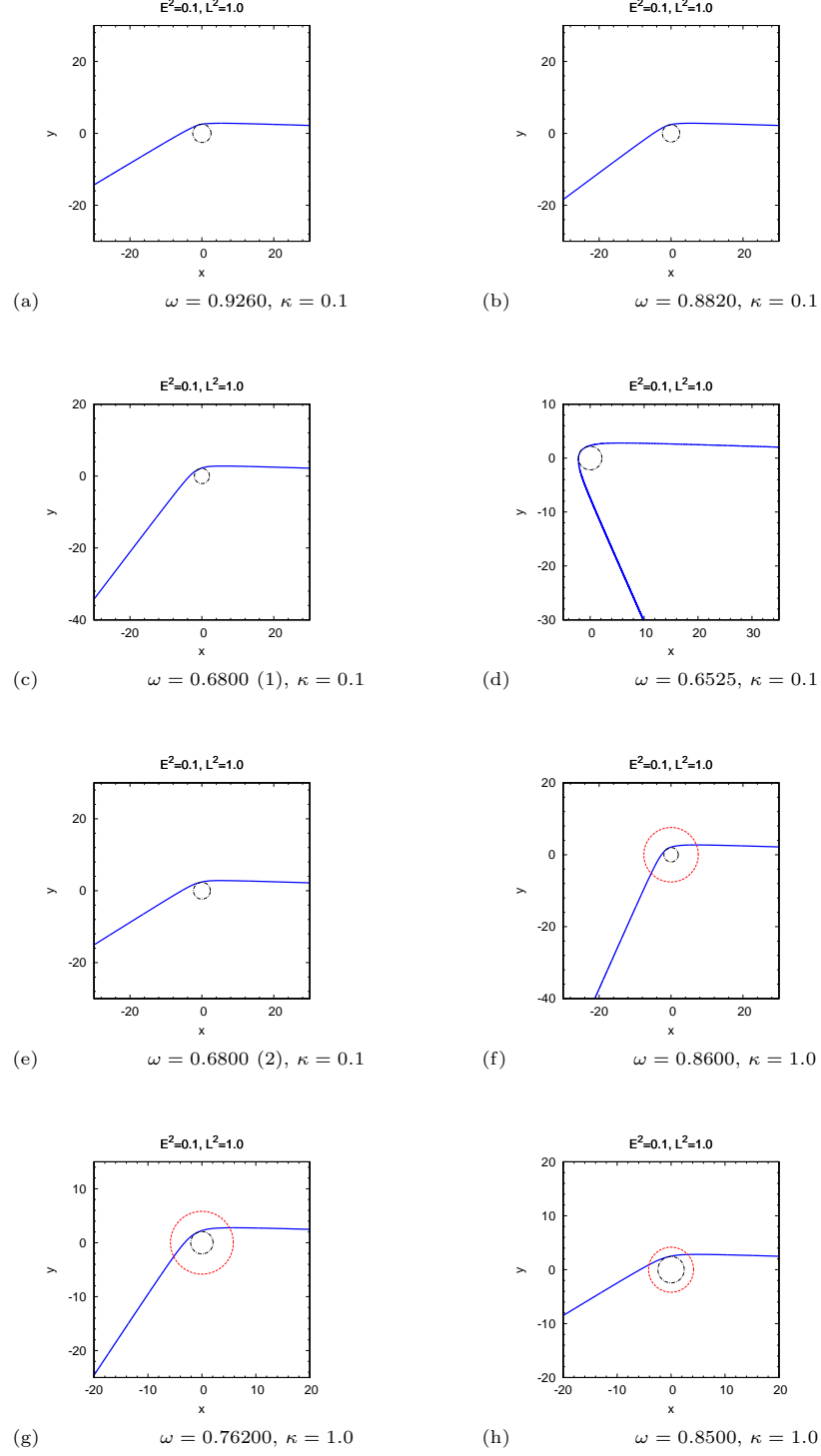


FIG. 6: We show escape orbits of massless test particle with $E^2 = 0.1$ and $L^2 = 1.0$ in the space-time of a boson star with $\kappa = 0.1$ and $\kappa = 1.0$, respectively for the boson stars given in Table I. Also indicated is the radius of the event horizon of a Schwarzschild black hole with the same mass as that of the boson star (red circle) if this radius is larger than the minimal radius of the escape orbit, which is the case only for $\kappa = 1.0$.

[38, 39], which is contained within a sphere of radius roughly $2 \cdot 10^{10}m$. Comparing with Table II we observe that we can have boson stars that can be as dense as the observed object. Furthermore, since the orbits of a number of stars have been determined, we can compare the observational results with ours. Unfortunately, the perihelion shift of the orbits of stars around the galactic center has not yet been determined (further observations are necessary here), but the eccentricity $e := \sqrt{(r_{\max}^2 - r_{\min}^2)/r_{\max}^2}$ of the orbits has been calculated from the observations. In Table V we give some values of e for orbits in boson star space-times. Note that these are the nearly maximally available eccentricities, but that we can also have orbits which are nearly circular (see discussion above on ISCOs). Hence $e \geq 0$ up to $e \lesssim 1$ is possible. Interestingly, some of the orbits of stars circling around the center of our own Milky Way have quite large eccentricities [40–42], e.g. the orbit of the star called *S14* could have eccentricity up to $e \approx 0.974 \pm 0.016$ [40]. Clearly, we are able to describe these large eccentricities with our model.

Another interesting observation has been done related to the emission of short wavelength radio waves from *Sagittarius A**. These are emitted due to accretion of surrounding matter onto the supermassive object in the center of our Milky Way. In [43] radio waves with wavelength of 1.3 mm have been observed in order to understand the region close to the supermassive object in which gravitational effects should be very strong. Observing *Sagittarius A** at wavelength 1.3 mm, the authors of [43] concluded that the size of this object is smaller than the expected apparent event horizon size of the assumed black hole. It was suggested that this might be due to the fact that the biggest fraction of emission of radio waves is not due to the black hole at the center, but rather related to the accretion in the vicinity of the supermassive object.

If we would, however, give up the idea of a black hole in the center of our galaxy and use a non-compact boson star instead particles could be emitted from radii smaller than the apparent event horizon of the black hole since the space-time is globally regular. This is demonstrated in Fig.7, where we show bound and escape orbits of massive test particles in the space-time of a boson star with $\kappa = 1.0$ and different values of ω . For comparison we also show a circle with radius equal to the Schwarzschild radius of a black hole with the same mass, which in our coordinates is $r_h = M/2$. It is apparent that the minimal radius of these orbits (both bound as well as escape) is smaller than the event horizon, while the maximal radius of the bound orbits is larger. If we would hence have stricter bounds on the mass (and hence the corresponding Schwarzschild radius) of the supermassive object in the center of our galaxy and would have clear signals from particles coming from “inside” the event horizon a possible explanation might be a non-compact boson star. Future observations at wavelength even smaller than 1.3 mm should be able to show what the space-time of the supermassive object at the center of our galaxy looks like.

2. Perihelion shift

The motion of a number of stars orbiting the center of our Milky Way have been determined [40–42]. Since the orbital periods T_o of these stars are quite large (between $T_o = 11.5 \pm 0.3$ years for S0-102 [42] and $T_o = 94.1 \pm 9.0$ years for S1 [41]) it will need longer observations to determine the perihelion shift of these objects. But this is certainly one of the interesting observables. Hence, we discuss the possible perihelion shifts of the orbits of massive test particles in the space-time of non-compact boson stars here. These might be later used to compare with observations.

The perihelion shift of a bound orbit of a massive test particle is given by

$$(\Delta\varphi)_{\delta=1} = 2 \int_{r_{\min}}^{r_{\max}} \frac{dr}{\sqrt{E^2 - V_{\text{eff}}(r)}} \frac{Lf}{r^2 \sqrt{l}} - 2\pi . \quad (33)$$

with V_{eff} given by (3), where we set $\delta = 1$.

The values for orbits with nearly maximal eccentricity are given in Table VI. We observe that the perihelion shift in all cases is quite large and can even become negative. This latter fact would be something that is not possible in a Schwarzschild black hole space-time and would allow us to distinguish black holes from boson stars.

3. Light deflection

Since the gravitational field close to a supermassive object is very strong it is important to understand light deflection in order to be able to interpret observational data correctly. Due to the strong bending of light rays the apparent event horizon of the proposed black hole would be twice as big as the actual event horizon. All of these things have to be taken into account when analyzing observational data.

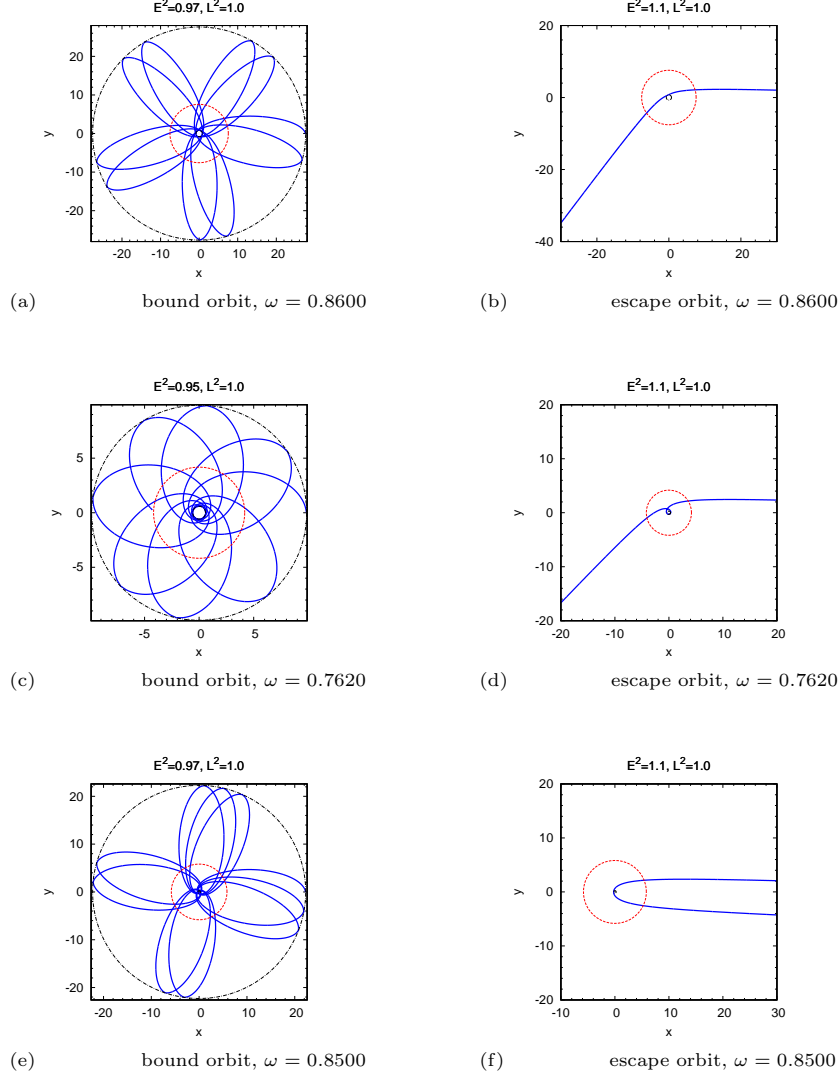


FIG. 7: We show bound and escape orbits of massive test particle in the space-time of a boson star with $\kappa = 1.0$ for different values of ω . Also indicated is the Schwarzschild radius $r_h = M/2$ (red circle), i.e. the radius of the event horizon of a black hole with the mass M equal to that of the respective boson star (see also Table I).

ω	κ	E^2	L^2	$(\Delta\varphi)_{\delta=1}$
0.9260	0.1	0.97	1.0	-0.5122π
0.8820	1.0	0.95	1.5	-0.1448π
0.6800 (1)	0.1	0.95	0.5	-0.2284π
0.6525	0.1	0.95	1.5	1.3566π
0.6800 (2)	0.1	0.95	0.5	0.4933π
0.8600	1.0	0.97	1.0	-0.4161π
0.7620	1.0	0.97	1.0	0.4852π
0.8500	1.0	0.95	1.0	1.2386π

TABLE VI: The values of the perihelion shift $(\Delta\varphi)_{\delta=1}$ for the boson star space-times given in Table I and for the orbits with nearly maximal eccentricity (see also Table V).

ω	κ	E^2	L^2	$(\Delta\varphi)_{\delta=0}$
0.9260	0.1	0.1	1.0	0.1619 π
0.8820	0.1	0.1	1.0	0.1936 π
0.6800 (1)	0.1	0.1	1.0	0.2829 π
0.6525	0.1	0.1	1.0	0.6228 π
0.6800 (2)	0.1	0.1	1.0	0.1678 π
0.8600	1.0	0.1	1.0	0.3568 π
0.7620	1.0	0.1	1.0	0.3039 π
0.8500	1.0	0.1	1.0	0.1636 π

TABLE VII: The values of the light deflection $(\Delta\varphi)_{\delta=0}$ for the boson star space-times given in Table I and $E^2 = 0.1$ and $L^2 = 1.0$.

The light deflection in the space-time of a boson star is given by

$$(\Delta\varphi)_{\delta=0} = 2 \int_{r_{\min}}^{\infty} \frac{dr}{\sqrt{E^2 - V_{\text{eff}}(r)}} \frac{L f}{r^2 \sqrt{l}} - \pi . \quad (34)$$

with V_{eff} given by (3), where we set $\delta = 0$.

In Table VII we give the light deflection in the boson star space-times indicated in Table I for massless test particles with $E^2 = 0.1$ and $L^2 = 1.0$. We observe that the light deflection increases when moving along the branch for $\kappa = 0.1$ up to roughly the minimal possible value of $\omega \approx 0.6525$. On the second branch, the light deflection decreases again. Similar to what was stated above for massive test particles, there seems to be no direct correlation between the mass of the boson star and the value of light deflection and hence the presence of the scalar field seems to have a non-trivial effect. For $\kappa = 1.0$ we find that the light deflection decreases when moving along the branches. This again indicates that a semi-analytical treatment of the space-time of boson stars does not lead to the proper results and that numerical techniques are vital to make predictions about observables in such kind of space-times.

V. CONCLUSIONS AND OUTLOOK

In this paper, we have studied the geodesic motion of test particles in boson star space-time and have discussed possibilities to differ them from the geodesics around a Schwarzschild black hole. Our work has many interesting applications. As a first example let us mention the accretion of matter onto massive compact objects such as neutron stars. Neutron stars are very dense objects that are typically modeled using an appropriate equation of state. Boson stars could act as a toy model for these objects. In the following, one could use our techniques to model accretion flow onto a boson star in a simple ballistic approach. This was done before in black hole space-times where analytic solutions to the geodesic equation exist [44, 45]. This technique could be applied as well using the numerical data of a boson star space-time. The accretion flow would also be interesting from the point of view of interpreting the boson star as an alternative to a supermassive black hole residing in the center of our Milky Way. The accretion of matter leads to the emission of radio waves and these can be used to make predictions about the properties of the supermassive object at the center of our galaxy. Gravitational wave signals would be another possibility. These have already been discussed in [30].

It would also be interesting to see how test particles move in radially excited boson star space-times as well as in the space-time of a rotating boson star. For once this is interesting since it is believed that the flat space-time counterparts of boson stars, so-called Q -balls could originally form in an excited state [6]. Furthermore, due to the interaction with surrounding matter, e.g. accretion, boson stars might be excited. If boson stars should act as toy model for neutron stars it would be surely also of interest to study the rotating counterparts since basically all observed neutron stars seem to rotate, in fact quite quickly as the example of a radio pulsar (assumed to be a neutron star) spinning at 716 Hz shows [46]. It would then be interesting to study the motion of test particles in these space-times.

Finally, in order to model accretion flow onto dense astrophysical objects more realistically one should consider charged test particles (that would make up the plasma). Hence the equation of motion for the particle should be modified to include the charge of the particle. Studies of charged test particles in some black hole space-times have been done previously [47]. In order to have interactions of the test particles with the boson star the latter should be charged as well. This can be done by promoting the global $U(1)$ symmetry of our model to a local $U(1)$ and has been considered for a number of boson stars [48]. The corresponding equation of motion for the test particle would then

have to be solved numerically.

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